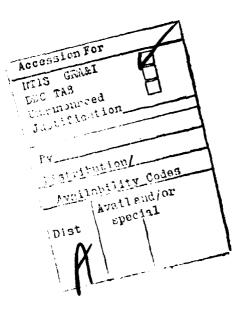


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Air Force. Conclusions are drawn concerning the identification of AFGL sources with IRC sources. A rough classification is made of the catalogue objects within our area of the sky. Limits for completeness of the survey are approximately +1.0 mag at 4 mm and -1.0 mag at 11 um. moromesers

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INTRODUCTION

This report summarizes the results of the UCSD group's ground-based, photometric observations of sources detected in the AFGL infrared sky survey. This report will be concerned principally with sources from the AFGL catalogue although the computer readouts enclosed with this report detail all UCSD observation, including those of many AFCRL objects excluded from the revised catalogue.

By agreement with the U. of Minnesota and U. of Wyoming groups, we observed the 40-60 mzone in each hour of right ascension. Practical declination limits were +60 (set by the yoke mount of the Mt Lemmon telescope) and -30. These boundaries encompass ~23% of the total sky. There were 616 AFGL sources within these limits of which 271 were scanned. This included 99 of the 106 unknown sources.

There are six sections following this introduction: the first reviews the observational procedures; the second summarizes the observations; the third compares our observed magnitudes with those of the Air Force; the fourth discusses Air Force identification with IRC objects; while the fifth gives a rough classification of 616 sources within the observational boundaries. The final section evaluates the limits of completeness of the AFGL Survey.

OBSERVATIONAL PROCEDURE

All observations were made with the 1.5 m, f/16 telescope at Mt. Lemmon. All sources, whether known or unknown, were scanned for by first locating a chosen SAO star and then offsetting to the SW corner of a box 4'.5 (EW) by 6'.0 (NS) centered on the AFGL coordinates. This box size represents ± 2 g in each coordinate. The error box was then raster-scanned for the source. The scan rate was 40"/sec. The aperture was 17" and the beam throw 26". A typical chopping frequency was 20-30 Hz. Upon completion of the raster, the telescope was again moved automatically to a second SAO star to serve as a check on the procedure. Scans were done primarily with a glass filter used with a Si(As) system. The scanning detection limit with this system was 5.0 mag at 3.5 µm. Some early scans were done with a InSb system with a detection limit of 7.0 mag at 3.5 µm.

If the source was detected, then broadband photometric data were obtained. Two different Si(As) systems were used with slightly different sets of broad band filters. Observations prior to May, 1978, were obtained with a downlooker system with the filters specified in Table 1, while measurements from May, 1978 were made with an uplooker system with filters shown in Table 2.

Some early observations with the InSb system included broadbands at 1.65, 2.3, and 3.5 μ m. In addition, some spectrophotometry at $\Delta \lambda/\lambda \sim 2\%$ was obtained from 2-4 μ m and 8-13 μ m for some of the more unusual sources. These are given in the enclosed A.J. paper by Gosnell et al. (See Appendix B)

Some very red objects may have been missed by scanning at 3.5 um. However, it is computed that with the UCSD system a detection is more likely scanning at 3.5 um than at 11.2 um for an object warmer than 410 K.

Since the beam separation was only 26", it is likely that many of the real sources that were not detected were extended. This is evidenced by the failure to detect HII regions.

TABLE 1

BROAD BAND FILTERS
Si(As) DOWNLOOKER

<u>λ(μm)</u>	Band Pass (um)
2.28	0.5
3.50	1.0
4.90	1.0
8.40	0.8
11.20	2.0
12.50	1.7

TABLE 2 BROAD BAND FILTERS

Si(As) UPLOOKER

$\lambda (um)$	Band Pass (um)
2.28	0.41
3. 50	1.00
4.74	0.50
8.48	0.81
10.55	0.90
11.94	1.04
12.52	1.20

OBSERVATIONS

The objects we observed do not represent an unbiased sample of our section of the sky. An attempt was made to observe as many unknowns as possible (of the 106 all but 7 were scanned). Also preferential weighting was given to very red objects, sources with unexpected magnitudes, stars without known spectral types, and stars with shells. It was hoped to obtain photometry on all known Mira variables in the zone. Observations of bright stars and well investigated sources (e.g. IRC +10 216) received less emphasis.

The sources scanned are classified under the following headings: normal stars, shell stars, late-type variables without indication of shells, peculiar stars without indications of shells, non-stellar sources, unknowns which were found, unknowns not found, knowns which were not found, and stellar sources which were difficult to classify. (The expression "known" refers to any source associated by the AFGL with some previously detected objects.

Sources classified as "not found" were those not detected in UCSD scans for which scanning procedures were correctly followed, i.e., the box scanned was correctly centered on the AFGL coordinates as indicated by the SAO stars.)

Normal stars are those which are not known variables and which rise less than 0.4 mag between 3.5 and 11 µm. Shell stars show rises of 0.6 mag or greater. Peculiar stars without indication of shells are variables other than late-type variables (e.g. R CrB) or IRC sources with I-K values > 3.5 but 3.5 to 11 µm rises of less than 0.4 mag.

There are two non-stellar sources associated with knowns. GL 4141 is identified as LHE 316. GL 4029 is incorrectly identified as LX Cas, an eclipsing variable, but is in fact a very red, multiple source which is being further investigated.

Three of the unknowns which were found, GL 2009, 2252, and 2287, were searched for but not located by other observers. One of the unknowns (GL 2999) which we were not able to find was located by another group.

Of the 11 knowns which were not detected in the scans, three (GL 998, 2046, and 2639) fall outside the $\pm 2\sigma$ error box. A third, GL 2304, was an extended source (W48).

Stellar sources which were not classified as either normal or shell are ones for which 10 um photometry were not obtained.

TABLE 3
OBSERVED NORMAL STARS

AFGL	AFGL	AFGL	AFGL
108	635	1293	2261
111	671	1372	2267
113	674	1378	2278
123	819	1379	2481
129	820	1387	2491
132	862	1438	2796
143	1001	1446	2828
262	1004	1509	2986
274	1191	1510	4007
276	1216	1583	4042
279	1218	1584	4139
377	1285	2235	4198
519	1289		- , -

TABLE 4
OBSERVED STARS WITH SHELLS

AFGL	AFGL	AFGL	AFGL	AFGL	AFGL
122	527	999	1434	1743	2452
127	528	1009	1437	1793	2479
253	622	1028	1439	1801	2645
257	639	1043	1441	1805	2818
272	664	1045	1450	1806	2819
278	667	1049	1511	1818	2940
280	812	1050	1516	1905	2941
287	815	1052	1519	2037	2974
371	832	1173	1576	2040	3143
378	837	1184	1579	2241	3194
379	842	1288	1594	2254	4013
381	846	1298	1650	2270	4065
414	850	1301	1652	2289	4241
521	853	1302	1660	2300	4269
522	856	1376	1669	2443	4295
525	858	1380	1726		

TABLE 5

LATE TYPE VARIABLES WITH NO SHELLS

AFGL	VARTABILITY
120	V451 Cas
1880	UV Her (Mira)
2251	AB Aql (LG)

TABLE 6

PECULIAR STARS WITH NO INDICATIONS OF SHELLS

AFG.
314
357
404
512
624
643
661
4219

TABLE 7

NON-STELLAR OBJECTS

AFGL

4029

4141.

TABLE 8

UNKNOWNS FOUND

AFGL	AFGL	AFGL	AFGL
538	2015	2287	2636
809	2023	2290	2679
1039	2047	2445	2686
1894	2252	2477	4253
2009	2259	2494	4306

TABLE 9

AFGL UNKNOWNS NOT FOUND

AFGL	AFGL	AFGL	AFGL	AFGL	AFGL
115	649	1027	1902	2681	4014
124	650	1033	1917	2824	4036
126	669	1188	2051	2939	4038
144	818	1292	2242	2961	4040
250	821	1295	2243	2999	4041
256	845	1392	2245	3008	4079
387	857	1444	2256	3139	4080
389	860	1453	2271	3144	4089
409	1002	1514	2303	3151	4184
412	1018	1523	2448	3178	4240
530	1020	1575	2455	4009	4242
621	1024	1659	2492	4010	4254
634	1026	1734	2670	4012	4257
645					

TABLE 10

KNOWNS NOT FOUND

AFGL
823
998
1017
1038
1900
2046
2266
2304
2639
3159
3185

TABLE 11

UNCLASSIFIED

AFGL
254
403
627
4138

A COMPARISON OF AFGL AND UCSD MAGNITUDES

The following two figures are histograms of the AFGL 4 µm and 11 µm magnitudes minus the UCSD magnitudes at roughly the same wavelength. The AFGL magnitudes are -0.40 brighter at 4 µm and -0.42 brighter at 11 µm. The dispersion of the curves about the mean value is 0.53 at 4 µm and 0.84 at 11 µm.

The UCSD 4.2 um magnitude was obtained by weighting the 3.5 um and either 4.7 um or 4.9 um fluxes inversely by their separation in wavelength from the AFGL filter. A careful correction for the CO in the 4.7 or 4.9 um bandpasses may produce better agreement between the two systems.

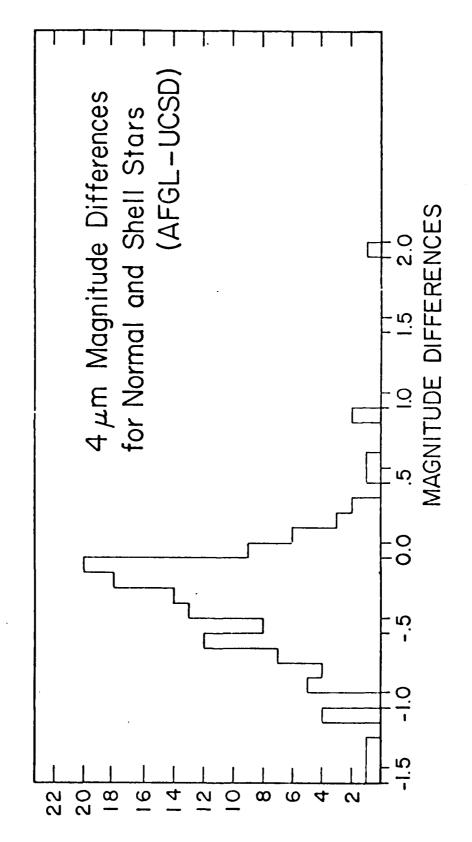
The UCSD 11.0 µm magnitude was obtained from either the 11.2 µm measurement or a simple mean of the 10.6, 11.9, and 12.5 µm magnitudes.

The histograms include all sources, regardless of their apparent magnitude. In order to assess the accuracy of the AFGL data and to aid in estimating the level to which the survey is relatively complete, scatter plots of the 4 um and 11 um data were formed. With each is a table presenting the mean and standard deviation of the differences of the AFGL and UCSD magnitudes as a function of AFGL magnitude. Five sources with highly discrepant 11 um magnitudes which were omitted from the 11 um plot are listed in Table 14. They appear to be either spurious detections or else are regions of extended emissions.

To gain some idea of the contribution to the scatter due to variability

of the sources, the IRC-UCSD 2.2 μm magnitudes we plotted against AFGL 4 μm and 11 μm magnitudes. Two highly variable objects, R Lyn and SU Mon, are included in the 4 μm plot but not in Table 15. Their inclusion increases σ to 0.26.

The 4 um plot includes sources detected only at 4 um by the AFGL. From the plots and Tables 15 and 16 the greater variability of the 11 um sources is apparent. Assuming the variations at longer wavelengths are no greater than at 2.2 um, the contributions to the scatter at 4 um and 11 um are no larger than 0.20 and 0.40 magnitudes, respectively. This leaves errors of 0.35 mag at 4 um and 0.45 mag at 11 um to be accounted by AFGL and UCSD measurement errors and by systematic differences such as beam size.



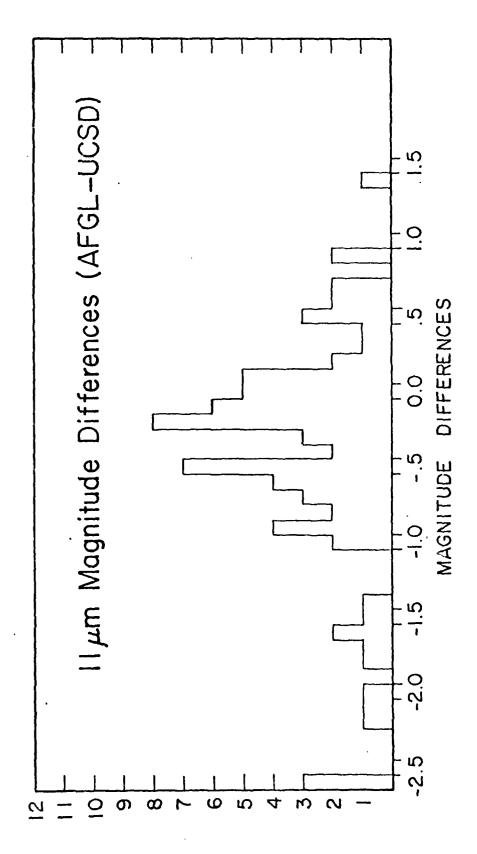


TABLE 12

AFGL 4 µm MAGNITUDES VERSUS AFGL (4 µm) - UCSD (3.5 µm) MAGNITUDES

AFGL mag	Number of Observations	Mean (AFGL-UCSD)	σ (AFGL -U CSD)
-1. 0	8	-0.22	0,36
- 0 . 5	11	-0.17	0.35
0.0	19	-0.16	0.35
0.5	27	-0.21	0.34
0.7	30	-0.24	0.33
0.8	34	- 0.27	0.34
0.9	39	-0.24	0.38
1.0	41	-0.27	0.40
1.1	45	-0.26	0.39
1.2	5 3	-0.28	0.43
1.3	57	- 0 . 29	0.44
1.5	68	-0.36	0.49



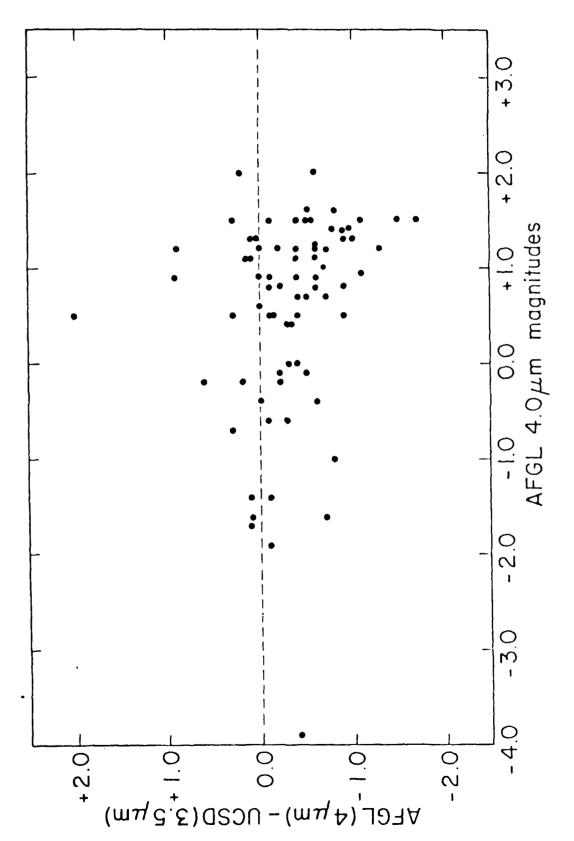


TABLE 13

AFGL 11 µm MAGNITUDES VERSUS AFGL (11 µm) - UCSD (11 µm) MAGNITUDES

AFGL Mag	Number of Observations	Mean (AFGL-UCSD)	a (AFCL-UCSD)
-2.4	12	0.22	0.49
_1.8	15	0.13	0.50
-1. 6	21	-0.16	0.74
-1. 5	24	-0.22	0.72
_1.4	27	-0.26	0.69
-1. 3	36	-0.30	0.61
_1.2	41	-0.25	0.61
-1.1	49	-0.33	0.65
-1.0	53	-0.32	0.65

AFGL #	AFGL (11 µm)	UCSD (11 μm)	ΑΡΟΙ. (4 μm) - UCSD (3.5 μm)	IRC (2.2 µm) - UCSD (2.28 µm)
262	-1.0	1.6	-0.5	0.03
1191	-1.8	0.8	-0.2	-0.15
1438	-2.1	0.1	-0.3	-0.06
1805	-2.1	0.5	- 0.5	0.00
2974	 1.5	0.5	-0.4	0.15

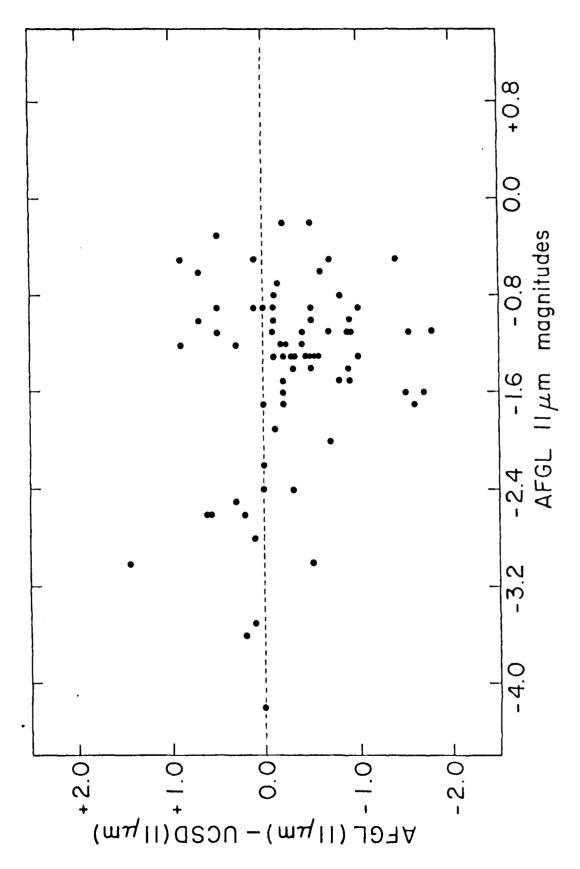


TABLE 15

AFGL 4 µm MAGNITUDES VERSUS IRC (2.2 µm) - UCSD (2.28 µm) MAGNITUDES

AFGL mag	Number of Observations	Mean (AFGL_UCSD)	a (AFGL-UCSD)
0.0	3	-0.08	0.07
0.5	7	-0.19	0.14
1.0	25	-0.12	0,20
1.5	58	- 0 . 08	0,18
2.0	63	-0. 08	0.18

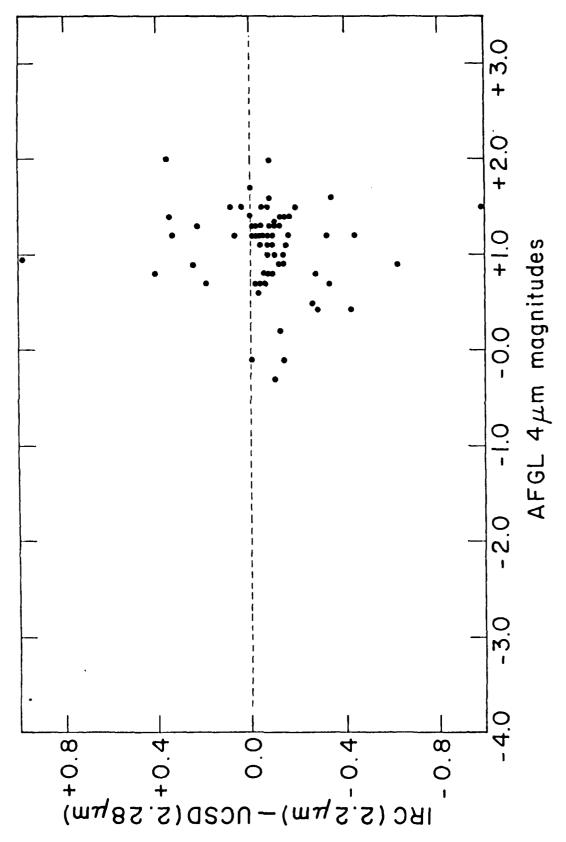
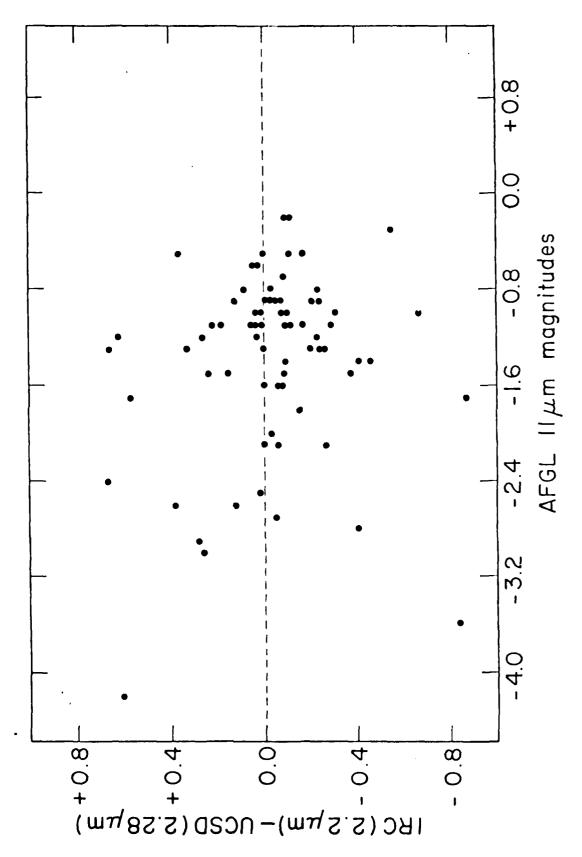


TABLE 16

AFGL 11 4m MAGNITUDES VERSUS IRC (2.2 4m) - UCSD (2.28 4m) MAGNITUDES

AFGL Mag	Number of Observations	Mean (AFGL-UCSD)	a; (AFGL-UCSD)
-2.4	11	- 0 _• 05	0.65
_1.8	16	-0.07	0.54
- 1.6	21	-0.07	0,52
-1.5	25	-0.06	0,48
-1.4	28	-0.09	0.47
-1.3	34	-0.07	0.45
-1. 2	38	-0.04	0.44
-1.1	47	-0.04	0.40
-1.0	53	-0.06	0.39
-0.9	60	-0.06	0.37



IDENTIFICATION WITH IRC SOURCES

We have attempted to check the identification of AFGL sources with IRC sources for the 166 such objects for which we have data.

Of the 166, 107 have K magnitudes within ±0.20 mag of the IRC values. These are taken to be correctly identified as the IRC objects. Twelve more objects with magnitude differences outside these limits, but which were identified as SAO stars are confirmed as the IRC sources by positional agreement.

There are 31 sources which agreed in magnitude to within ± 0.30 mag of the IRC, or were brighter than the IRC values. These are designated as probable identifications.

The remaining 16 objects did not satisfy any of the above criteria. However, many of these are late-type variables and two are noted as variables in the IRC. There were no sources for which the AF identification with IRC objects were found to be in obvious error.

TABLE 17

AFGL OBJECTS CORRECTLY IDENTIFIED WITH IRC OBJECTS

AFGL	AFGL	AFGL	AFGL	AFGL	AFGL
108	371	674	1285	1509	2267
111	377	812	1288	1511	2278
113	378	819	1289	1519	2452
120	379	820	1293	1576	2479
123	381	832	1298	1583	2796
127	403	858	1302	1594	2818
129	414	862	1372	1650	2828
132	512	873	1378	1652	2932
143	519	999	1379	1669	2940
253	525	1001	1 380	1726	2941
254	528	1004	1387	1801	2974
257	622	1009	1434	1805	3196
262	624	1043	1437	1889	4007
272	627	1045	1438	1905	4013
274	635	1184	1439	2037	4042
276	639	1191	1441	2235	4198
278	643	1216	1446	2254	4269
279	671	1218	1450	2261	

TABLE 18

AFGL OBJECTS CORRECTLY IDENTIFIED WITH IRC

OBJECTS BY POSITIONAL CORRELATION

AFGL	SAO
280	022817
517	024202
1050	026097
1173	153273
1510	043886
1511	043889
1660	182081
2481	143959
2491	069246
2819	164760
3140	053335
4139	062754

TABLE 19

AFGL IDENTIFICATIONS WITH IRC OBJECTS PROBABLE

AFGL	AFGL	AFGL	AFGL
122	856	1660	2645
287	1028	1743	2819
517	1131	1818	2986
664	1173	2241	4138
667	1376	2251	4241
837	1510	2289	4269
842	1516	2443	4295
846	1579	2491	

TABLE 20

AFGL IDENTIFICATION WITH IRC NOT CERTAIN

AFGL	Variability
404	
521	
527	(K)
815	
850	
853	AZ Aur (Mira)
1050	R Lyn (Mira)
1052	RS Gem (SRb)
1793	BG Ser (Mira)
1806	Eu Lib (SR)
2040	T Dra (Mira)
2270	V2059 Sgr (Mira)
2300	ST Sgr (Mira)
2993	S Aqr (Mira) (K,I)
3143	EY And (Mira)
4065	X CMa (SRb)

BREAKDOWN OF OBJECTS IN OUR ZONE

As mentioned previously, the 271 objects we scanned for do not constitute an unbiased sample of our area of the sky. However, by using the AFGL magnitudes and identifications, I-K values from the IRC, and variability type a reasonably accurate breakdown can be made.

When this is done approximately 43% of the objects are found to be stars with shells and 37% to be stars without shells. Non-stellar knowns, mostly HII region, constitute 3%. Unidentified sources complete the total with 17%. Of the unidentified sources, 20% have been located through ground-based observations. A substantitally higher fraction could nevertheless be real if they are extended sources.

Of the unknowns detected, several are very red objects. Except for GL2659 and 2287, the preliminary observations of the sources have been reported by Gosnell et al.

THE LIMITS OF THE AFGL SURVEY

Unfortunately, the scatter plots of the previous section do not clearly determine the levels of statistical completeness of the surveys at 4 um and 11 um. To supplement these plots, this section presents tables of source counts versus apparent magnitude at the two wavelengths. Since the infrared data are not as severely affected by interstellar reddening as the optical and since the survey limits are not faint enough to see substantial decreases when looking towards high galactic latitude, the source number "N" versus apparent magnitude "m" should approximate a log N = 0.6m + Constant relationship. The source counts, taken from UCSD data, are tabulated in half magnitude increments. In fact, the slope is something less than 0.6 at both wavelengths. Despite this, and although the data is subject to uncertainties due to the number of sources, the 3.5 um data clearly flattens out beyond +1.0 magnitude, while the 11 um data turns over -0.5 magnitude. Combining these data with the scatter plots of the previous section yields a best estimate for the survey limits as +1.0 magnitude at 4 um and -1.0 magnitude at 11 um, subject to errors of a few tenths of a magnitude.

TABLE 21

3.5 µm SOURCE COUNTS

UCSD Mag	N	Log N	Δ(Log N)
-3. 5	1		
- 3.0	1		
-2.5	0		
-2. 0	0		
-1. 5	4		
-1.0	4	0,60	0.0
- 0 , 5	7	0.85	0.15
0.0	11	1.04	0.19
0.5	19	1,28	0.24
1.0	32	1.51	0.23
1.5	35	1.54	0.03
2.0	3'7	1.57	0.03
2.5	24	1.38	-0.19
3 . 0	5	0.70	-0.68

TABLE 22

11 um SOURCE COUNTS

UCSD Mag	N	Log N	Δ(Log N)
-4.5	1		
-4. 0	2		
-3. 5	2		
-3. 0	4		
-2.5	4		
-2.0	4	0.60	0.00
-1.5	6	0.78	0.18
_1.0	12	1.08	0.30
-0.5	20	1.30	0,22
0.0	10	1.00	~ 0.30
0.5	7	0.85	-0.15

UCSD PERSONNEL WHO PARTICIPATED IN THE AFGL PROGRAM

The program of photometry of AFGL sources at UCSD has been contributed to by a large number of people. Observations were made by A. H. Hewitt, Roc Cutri, Timothy R. Gosnell, Hugh S. Hudson, Paul D. LeVan, C. A. Lindsey, R. C. Puetter, Richard J. Rudy, B. T. Soifer, and S. P. Willner. Technical support was provided by Paul Brissenden, John Casler, Dawn Pedersen, Eileen Smith, Pat McCune, and R. W. Russell. Shane Burns aided in the data reduction.

APPENDIX A

This appendix presents computer readouts of all the UCSD observations. The photometry for each object is presented along with the date or dates of observation and possible identification of the source as an SAO star. Various comments are also included; a code for these comments is given on the following page. The observations are classified by which system they were taken with (downlooker or uplooker) since the filters differ slightly between systems as mentioned in the observational section.

		UCSU UBSERVATIONS	* * * * *	
		GL/CRL SURVEY SOURCES		
• • •				
h-7	AND ONA	DOWNLOUKER UDSERVATIONS AS OF 05/13/79		

* * * * *			* * *	
	FERREST CUORS			
••••			***	
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****	• 3	EXTENDED UBJECT (MITH RESPECT TO THE CHE	***	
		BEAN UICE)	****	
	2		***	1
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••••	•	THECH UNE' EIC. (IE. HEAM SEPARATION & UNE J	***	٠ 1
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P. 0.24	•	09/24/78	۲ ۲ - ۲ ۱۵ - ۱۵	2.59	2.40	2.40	•	2.70	2.04	• •	• •	124	
642		- 01/02/60	2	2.2	2.3	2.5	• •	1.9	! • •	; • • !	• •	124	
500		02/18/79	L	1.12	0.80	0.52	•	•	02.0-	*0.74	-0.77	125	
0.0	11 52 430 640 750 71		\\\\\\\\\\\\\\\\\\\\\\\\\\\\\\\\\\\\\\	2.36	2.62	•	2.60	!	•	!	 -	120	
*			: YE.	3.23	3.11	2.94	•	3.25	•	•	•	25	
4178		05/21/10 05/30/76	X	1.37	1.56	1.87	• •	! • •		• •	: • •	130	
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APPENDIX B

This is a reprint of the paper published in the Astronomical Journal. It discusses the observations obtained with the downlooker system and is complete through May of 1978.

GROUND-BASED OBSERVATIONS OF SOURCES IN THE AFGL INFRARED SKY SURVEY

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ABSTRACT

We present a preliminary report on ground-based observations of sources discovered by the AFGL infrared sky survey. This paper contains photometric data on 104 sources and spectrophotometry ($\Delta\lambda/\lambda \approx 0.02$) of 14 sources identified in 6.0×4.5 -arcmin raster scans around the positions reported by the AFGL. The sources originally identified in the catalog with known objects are mainly hot, bright stars. Two unusual sources discovered in this program, GL 2636 and GL 4029, appear to be multiple sources. Both these objects exhibit the 3.3- μ m emission feature, and GL 4029 shows several other emission features, previously found in NGC 7027.

I. INTRODUCTION

When the preliminary version of the AFCRL infrared sky survey (Walker and Price 1975) was released in 1972, there was great hope that new classes of infrared objects would be discovered, especially among those sources not identified with normal stars. Subsequent ground-based observations, however, have shown that a majority of the sources not identified by Walker and Price are, in fact, late-type stars. Nevertheless, these observations (Cohen 1975; Cohen and Kuhi 1976, 1977; Lebofsky and Kleinmann 1976; Low et al. 1976; Merrill and Stein 1976; Gehrz and Hackwell 1976; Joyce et al. 1977; Lebofsky et al. 1978) have occasionally turned up remarkable objects.

Both to continue this search for new kinds of sources and to assess the catalog and its contents in an objective manner, we have initiated a systematic program of observations. These observations cover the fields of $\sim 1/3$ of the sources in the catalog (Price and Walker 1976), namely those lying between $40^{\rm m}$ and $60^{\rm m}$ of each hour of right ascension in declinations accessible from Mt. Lemmon, Arizona (2776 m, $07^{\rm h}$ $23^{\rm m}$ $09^{\rm s}$ 8 W, $32^{\rm o}$ 26' 21'' N).

II. OBSERVATIONS

All observations reported here have been obtained at the UCSD/University of Minnesota f/16 1.5-m telescope at Mt. Lemmon. The basic plan of the observations consisted of a standard raster scan of an error box centered on a catalog position. At the conclusion of the raster, any source discovered in the box was observed photometrically, and if the source proved interesting, narrow-band spectrometry was obtained. All classes of sources reported in the catalog have been included among the objects studied.

The standard raster pattern covered an area of 6.0 in hour angle, by 4.5 in declination, corresponding to a $\pm 2\sigma$ error box, where σ was determined from the systematic

positional errors reported in the AFGL survey. The telescope was first centered on an SAO star and slewed under digital control to the SW corner of the error box. Scanning then proceeded at a 0.7 s⁻¹ scanning speed and with a 17"-diameter circular aperture. North south chopping was used at frequencies of 10, 15, 20, and 30 Hz, and with a beam separation of 26". An east west line separation of 9" guaranteed that a source would appear in two consecutive lines, with the characteristic S-shaped form caused by the chopping making identification straightforward on the strip-chart recording. Finally, upon completion of the scan, the telescope again was slewed digitally from the NE corner of the pattern to a second SAO star. This procedure served as a check that all of the offset calculations had been performed accurately and that the correct part of the sky had been scanned. Closure errors varied from observation to observation, but were required to be smaller than 36".

Two detector systems were employed during the observations: An InSb photovoltaic detector was used for scanning at 3.5 µm, for obtaining 1.65 · 3.5 · µm broadband photometric data, and for obtaining 2 4 · µm spectrophotometric data. The second (long wavelength) system employed a Si(As) photoconductive detector, and was used for scanning both at 3.5 and 11.2 µm, for obtaining 2.3 · 12.5 · µm photometric data, and for obtaining 8-14 · µm spectrometric data.

The detection limits during scanning at $3.5~\mu m$ were 7.0 and 5.0 mag for the InSb and Si(As) systems, respectively. At 11.2 μm , the limit was 1.5 mag. Table I summarizes both the UCSD and the AFGL detection

TABLE UCSD survey sensitivities compared to those of the AFGL.

Wavelength (µm)	AFGI (mag)	UCSD (mag)
15		7 (5), InSb (St As)
4.2	+1 3	
11	-11	1.5. St As
19	- 3 1	

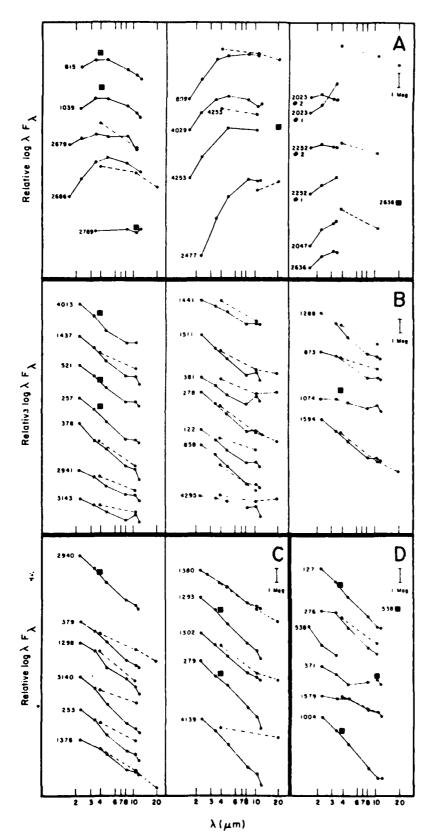


FIG. 1. Photometric data on sources in the AFGL infrared cotalog UCSD data are represented by solid lines, and AFGL data by dashed lines and squares. Only relative of use are shown, with AFGL data plotted at correct level with respect to UCSD data for a particular object. Numerical details appear in Table 1. (a) Cool objects ($T < 1000~\rm K)$, (b) objects with strong silicate or SiC emission at 11.2 μ m; (c) weak emission at 11.2 μ m; (d) unusual objects.

TABLE	H. Summary of ob	
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0108	00 43 55.7	+15 12 12 S		092072 K		:	0.29	•	0.05	0.27	0.10	-0.03	-0.03	ž
0111		+07 18 48 S		109474 K			0.89	•	0.66	0.88	0.83	0.66	0.80	3
0111	00 46 18.9	+56 48 10 S	10/15/76	021738 K		•	2.30	•	2.03	2.18	2.05	1.98	•	4
0115 0119	•		10/12/76	NF UI		•	•	•	•	•	•	•	•	5 6
0120	00 49 01.8	+59 18 06 S		021779 K		:	1,91	:	1.58	1.73	1.58	1.46	:	7
0122	00 49 53.	+47 08 .6 V		F K		•	2.21		1.62	1.43		-0.21		8
0123	00 50 27.0	-01 24 56 S		129009 K	_	•	1.06	•	0.85	1.09	0.99	0.94	0.85	9
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0126			10/15/76	NP U	K								•	11
0127		+48 24 29 S	_,	036763 K		•	1.77		1.54	0.69		-0.38		12
0129		+24 17 12 5		074365 K		•	1. 13	•	0.87	1.10	0.91	0.77	0.83	13
0132		+57 43 35 S -01 55 40 S		021846 K 129076 K		•	3.01 1.67	•	2.88 1.42	3.04 1.64	3.06 1.48	3.03 1.35	3.49 1.34	14 15
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0145	•		10/12/76	NP U		•	•	•	•	•	•	•	•	17
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0262 0262	01 51 41-	+08 32 .0 1	11/18/76	P K P K		2.18	1.99	•	1.85 1.86	2. 17	1.84	1.62	•	22 23
0274	01 54 52.9	+27 33 43 5		075048 K		1.71	1.50	:	1.33				:	24
0274	•		11/18/76	075048 K		•	1.51		1.29	1.64	1.33	1.33	•	25
0276	01 55 10.7	+39 53 31 5		055147 R		•	0.33	•	0.44		• • • • • • • • • • • • • • • • • • • •	•	•	26
0276 0278	01 66 27 3	+45 11 32 5	11/20/76	055147 K		•	1.00	•	-0.09 -1.35		-0.01 -1.88		- 2 90	27 28
02/8		-07 54 32 5		129546 K		•	2.15	:	2.17	1,68	1.86	1.75	2.37	29
0280		+54 34 49 5		022817 K			1.07		0.61		-0.06		-0.57	30
									2 10	2 20			0.40	
0371		+36 02 .3 N +29 02 27 S		P K 075578 K		•	2.69	•	2.38 1.90	2,20	1.94	-0.79 1.83	1.94	31 32
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0 181	02 46 55.3	+56 46 38 3		023647 K		•	2.10	•	1.48	1, 24	0.31	-1.26	-1.21	35
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0409	•		10/11/76		1K	:	:	÷	:	:	•	•	:	36
0412			10/11/76	NP U	JK		•	•					•	39
0416	•		10/25/77	P K		11.38	9.88	•	•	•	•	•	•	40
0416	•		10/24/77	FF		•	•	•	10.41	•	•	3.82	•	41 42
0443 0512	01 40 11.9	+12 18 11 1	12/08/76 5 12/06/76	NF 11 093566 K) K	•	1,12	:	0.79	0.97	0.74	0.53	0.48	43
0519		-12 15 26		149158 K		÷	0.37	:	0.13	0.27	0.16	0.07	0.08	44
0521		+50 41 32		024237 K		•	1.86	•	1.48	1.43		-0.11	0.02	45
0522	03 45 52.			F K		•	2.72	•	1.99 0.57	1.72	1.30	0.82	0.61	46 47
0525 0530	03 46 20.6	-07 10 00	10/11/76) K	•	0.88	•	0.57	0.04	0.42	0.10	-0.02	48
0538	03 58 00.5	+56 56 20			JK	5.25	5.80	•	5.45		•	•	•	49
0518			11/20/76	NP U) K	•	•	•	•	•	•	•	•	50
0621			10/11/76	NF L	υĸ									51
0622	04 40 59.	+20 40 .H			K	:	2.17	:	1.71	1.57	1.26	0.78	0.72	52
0624	04 40 34.0	+32 46 24	5 12/05/76	057393 H			0.97		0.68	0.65	0.61	0.56	0.53	53
66 14	•		10/11/76		UK	•		•	• • • •	1.33		1.21	1.32	54 55
0635 Ut 13	04 48 23.	+37 24 07 +28 26 .6			K K	•	1,53	•	1.26	0.76	1.31	0.02	0.04	56
0645	•	120 20 .0	10/11/76		O.K	÷	•	:	•	•	•	•	,	57
0643			10/11/76		UK			•	•	•		•	•	58
0650	•		10/11/76		UK UK	•	•	•	•	•	•	•	•	59 60
0649	•		10/16/76	ar c	U N	•	•	•	•	•	•	•	•	00
0671	34 54 57.	+hO 23	8 12/07/76	F 1	K		2.15		1,97	1.90		1.87		61
0674	04 58 59.		B 12/07/76		K		0.22	•	-0.05		0.00	-0.04	-0.02	62
9464 9464	05 40 31.	3 +12 40 58	10/11//6	_	UK NB	5.6A 6.63	3.3A 4.17	•	0.8A 1.67	•	•	•	•	63 64
0404	:		11/14/76		SK MB	•••	4.46	:	1,67	0.33	-1.50	-2.10	:	65
0 + 12	05 42 01.	7 +24 24 01		077502	K	•	1.83		1,11		0.00		0.00	66
0 4 1 5	05 44 0 %.	+43 11 .6			K	•	3.23	•	1.47		-0.48			67
0#15 0#1#	•		12/15/76		NK K N9	•	3.23	•	1.00	0.47	-0.44	-0.90		68 69
0+19.	05 44 55.5	5 -12 49 1R		150408		:	1.61	:	1.29	1.47	1,35	1.13	1.18	70
0#21 0#23	•		10/12/76 01/10/77		UK K	•	•	•	•	•	•	•	•	71 72
0-23	:		10/12/76		ELK.	:	:	:	:	:	:	:	:	73
0447			10/12/76	NF	UR			•	•			•		74
0 H S F	05 54 51.	•10 54 .8	V 12/06/76		K UK	•	1.33	•	0.47	0.70	0.37	-0.06	-0.26	75 76
0424	05 59 76.	-02 21 .2	10/12/76 ¥ 12/06/76		yr. H		-0.08	:	-0.41	-0.47	-0.A1	-1.65	- 1.59	77
0440			10/12/76	N.P.	UK	•		•						78
0862		3 +50 16 53			K	•	2.41	•	2.05	2.20		1.99		73
3 A 7 3	ወቀ ወያ ኃላ.	-05 42 .H	£ 02/03/17	P .	×		2.67		11.73	1.17	0.75	-0.15	-0.29	80

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73								02/21/78	F	K NV	11.44	8.71	•	5.00	•				н
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04	06	41	35.4	+29	01	24	S	12/05/76	078692			2,12	•	1.94	2.03	2.01	1.88	1,65	н-
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11			•					12/10/76	NP	UK	·	÷	÷	:	:	:	:	:	R
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18								11/18/76	NF	UK						•	•	•	
19			:					12/10/76	N P	UK	:	:	:	:	:	•	:	:	3
20 23			•					11/1H/76 02/0H/77	N F N P	UK UK	•	•	•	•	•	•	•	•	9
4			:					02/07/77	NE	UK	÷	÷	:	:	:	:	:	:	9
?6 ?7			•					02/07/77	NF	UK	•	•	•	•	-	•	•		9
9			:					02/01/11	N F N F	UK	• •	:	:	:	:	:	:	•	9
) 1) 1			•					02/08/17	NF	UK	•	•		•	-	•			9
			•					02/22/18	N P	UK	•	•	•	•	•	•	•	•	10
3			•					02/08/77	N P N P	OK OK	•	•	•	•	•	•	•	•	10 10
9					_			02/06/11	N P	UK	•	•			•		:	:	10
9	06	53	09.7	-02	16	18	U	02/07/77	P NP	UK	•	4.37	•	2.46	1.46	0.31	-0.03	0.05	10
8			÷					02/22/78	n P	UK	:	:	:	:	:	:	•	:	10
9			٠					02/22/78	N P N P	UK	•	•	•	•	•	•	•	•	10
4	07	05	26.	- 10	39	. 5	I	02/03/77	F	K	:	3.73	:	2.43	1.73	0.42	-0.42	-0.52	10
1	07	27	01.	-19	21	. 4	1	02/03/77	F	K	•	2.08	•	0.81	0.20	-0.82	-1.33	- 1. 25	1 1
0			•					02/08/77	NF	UK	•		•					•	1 1
2 B			•					02/07/17	N F	UK UK	:	•	•	•	•	•	•	•	11
8			•					02/22/78	NP	UK				:	÷	÷	÷	:	1 5
9 1	07	44	17.1	+11	12	25	s	02/07/77 02/12/77	NP 060340	K UK	•	1.20	•	0.97	1.04	0.94	0.87	0.84	1:
2	•		•		•	• 1	•	02/08/77	NP	ĸ	:		:	•		•	•	•	
3			•					02/08/17 02/07/71	N F N F	UK UK	•	:	•	•	•	•	•	•	11
1			:					02/01/17	NP	UK	·	:	:	:	:	:	:	•	12
6 8			40.7					12/06/76	135380		•	1.27	•	1.03	1.09	0.99	1.00	0.98	12
5	08	41	50.7	+ 18	20	22	S	12/05/76	098087		:	1.53	:	1.39	1.42	1.31	1.46	1.20	12
8								12/06/76	117103		•	0.24 1.57	•	-0.11	-0.08	-0.51		-1.31	12
ó	•••		•		30	12	3	12/07/76	NP	ÜK	:	•	:			1, 32	1.41	1.31	12
2 3	0.0	45	54.7	412	47	٥,٥		12/07/76 12/05/76	NF 098143	UK	•	1.91	•	1.67	1.65	• • • •	•	•	12
ś	••	٠,	•	* 12	٠,	,,	•	12/05/76	88	ūκ	:	•	:	•	•	1.65	1.44	1.68	12
7			•					02/08/77	NP	OK	•	•	•	•	•	•	•	•	1
8								02/12/77	098230			0.39		-0.14			-0.87		1
1 2								02/06/77	080524 098266		•	1.03	•	0.25			-0.75 -0.88		1
5								02/08/77	NP	UK	:		:		•	•	•	-0.93	1
2	09	41	00.6	+14	15	05	5	12/08/76	098733 NP	K UK	•	1.20	•	0.96	1.09	0.97	0.95	0.96	13
5			. :					12/07/76	NP	UK	:	:	:	:	:	:	:	:	13
7	09	42	34.7	+34	44	34	S	12/05/76 12/07/76	061669 NP	R UK	•	-0.21	•	-0.97	-1.43	-1.95	-2.12	-2.84	13
8	09	43	00.1	+57	21	32	5	12/08/76	027377		:	0.33	:	0.02	0.15	0.03	-0.03	-0.02	14
9	09	43	31.8	+06	56	25	5	12/06/76	117898	K		1.67		1.46	1.62	1.46	1. 35	1, 31	14
0								12/05/76	098769	K	•	-2.47	•				-4.41		1 4
3			:					12/06/76 02/07/77	N P N P	UK	:	:	:	:	:	:	:	•	14
4			•					02/08/77	NF	UR	•	•	•	•	•	•	•	•	1 4
7	09	51	05.4	+06	11	41	s	02/07/77	NP 117975	UK K.	1.48	1.47	1.18	1.05	:	•	•		14
0			•				-	03/10/77	NP	UK	•	•	•	•	•	:	:		1 4
1			•					02/07/77 12/07/76	N F N P	UK UK	:	•	:		•	:	:	:	14
.1	10	46	10.	•0.0	46	. 4		02/08/17	ř	ĸ		1,82		1,47	1.34	, ,	0.11	0.13	11
1			59.					02/08/11	7	ĸ	:	2,61	:	1.85	1.26		-0.36		1'
4			,,'-					02/07/77	NP	UK	•	•	•				•		1 '
6	10	7.5	43.7	•••	27	09	2	02/08/17	118576 NP	R UR	•	-0.71	:	-0.95	-0.89	-1.00	-1.16	-1.19	10
7							_	02/07/77	NP	UK	•	•	•	•		·	•	:	1 '
9								05/31/78 05/31/78	119035 043886		•	0.10 1.06	•	0.03	0.21 1.02	•	•	•	1.
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GLO	POSITION (19	(SO) CATE (UT)	SAO	COHE	1,65	2.28	3.2	3.5	4.9	8.4	11.2	12.5	CRD
1524	•	02/08/77	NF	OK		•		•	•	•	•	•	165
1528 1529	•	02/06/77 02/06/77	n? N?	UK UK	•	:	•	:	:	:	:	:	166 167
1530	:	06/09/16	HZ	UK		•	•	•	•	•	•	•	168
1573 1574	•	02/06/77 02/08/77	n P N P	DK UK	:	:	:	:	:	:	:	:	169 170
	•	•			-	-	_						
1575 1576	12 42 47. +45	02/07/77 06/07/76 8 2 . 8	HP P	k Ok	:	:	:	:	:	:	:	:	171 172
1577	•	02/07/77	HP	UK	•	•	•	•	•	•	•	•	173
1578 1579	12 44 46. +04	02/07/77 02/12/77 ¥ 02/12/77	ny P	UK K	•	1.81	:	0.74	-0.20	-1.11	-1.78	-1.70	174 175
1580	•	03/10/77	NP	UK	•	•		•	•	•	•	•	176 177
1582 1583	12 51 4509	02/07/77 16 B 05/31/78	NP P	UK K	:	0.52	:	0.26	0.57	:	:	:	178
1587	•	02/08/77	NP	DK	•	•	•	•	•	•	•	•	179 180
1590	•	02/08/77	HP	UK	•	•	•	•	•	•	•	•	
1594	13 00 06. +05	3 27 .2 7 02/12/77	7 N7	K OK	•	-0.99	•	-1.37	-1.55	-1.95	-2.79	-2.90	181 182
1644 1645	:	03/12/77 06/12/76	NF	UK	:	:	:	:	:	:	:		183
1646	•	06/12/76	M P M P	OK Ur	•	•	•	•	•	•	•	•	184 185
1647 1649	•	06/10/76 06/07/76	NP	DK	:	:	:	:	:	:			186
1655	•	02/22/78	NP	D K D K	•	•	•	•	•	•	•	•	187 188
1657 1659	•	02/22/78 06/11/76	H P	UK	:	:	:	:	:	:	:	:	189
1660	13 52 29.9 -26	6 11 13 5 05/30/78	18208	1 K	•	1.75	•	1.31	1.36	0.97	•	0.32	190
1662	•	03/12/77	N P	UK	•								191
1664	•	06/10/76 06/07/76	N? N?	OK	•	•	:	:	:	•	:	:	192 193
1665 1667	:	02/22/78	HP	UK	:	:	•			•	•	•	194
1668	11 50 16 7 . 16	06/11/76 9 16 17 5 06/11/76	NP 06388	DK 1 DK	3.86	3.73	3.62	3.62	:	:	:	:	195 196
1671 1672	*	03/12/77	N.F	OK	•	•	•	•	÷		•	•	197
1674 1723	•	02/22/78 06/11/76	N P N P	U K U K	•	•	•	:	:	:	:	:	198 199
1725	•	06/07/76	NP	UK BO			:	•	•		•	•	200
1727	•	06/11/76 06/11/76	N P N F	UK	•	•	•	•	•	•	•	•	201 202
172+ 1730	:	05/29/78	NF	UK	:	:	:	•	:			•	203
1711 1735	•	06/11/76 06/07/76	N P H P	UK UK	•	•	•	•	•	•	•	•	204 205
1717	:	06/12/76	NF	UK	:	:	:	:	:	:	•	•	206
1738 1739	•	06/12/76 06/10/76	N P N P	UK	•	•	•	•	:	:	:	:	207 208
1742	:	02/22/78	NY	UK	÷		:		•			•••	209
1743	14 55 02.6 -1	2 14 15 S 05/30/78	1589	29 K	•	0.40	•	-0.07	0.00	-0.37	•	-1.01	210
1795	•	05/24/78	N F	UK	•	•	•	٠.,	٠.,	5 -0.96		- 1.76	211 212
1 H O 2	15 48 23.2 +1	15 17 02 5 05/30/7H 06/07/76	1017 NP	71 K UK	:	0.48	:	-0.31	-0.5	0.96	•	-1.76	213
1808	•	06/12/76	NF	UK	•	•		•	•	•	•	•	214 215
1812 1813	•	05/29/78 06/07/76	N P N F	UK UK	:	:	:	:	:	:	:	:	216
1815	•	06/07/76	N P	UK	•	•	•	•	•	•	•	•	217 218
1884 1897	•	06/07/76 06/12/76	N P N F	UK EC	, .	:	:	:	•	:	:	:	219
1901	•	06/12/76	NP	UK	•	•	•	•	•	•	•	•	220
1902	•	06/12/76		UK EC					•		•	•	221
1922	17 04 532	24 39 .0 A 06/07/76 06/07/76	7	UK UK BO		•	•	•	•	•	•	:	222 223
2009 2009	17 45 45.5 -2	28 53 00 S 06/09/76	1857	88 UK	5.09					:	:	•	224
2009 2009	•	06/10/76 06/11/76	1857	80 UK NE						:	:	:	225 226
2015	17 47 292	27 51 .2 A 06/07/16	· F	UK	•	•	•	•	•	•	•	•	227 228
2022	17 51 152	06/12/76 06/12/76 A 06/07/76			•	:	:	:	:	:	:	•	229
2023		06/12/76			1 8.1/	6.64		3.54	•	•	•	•	230
2023		06/12/76		UK S	2 6.8	5.64	5.11	4.91					231
2047		17 44 34 0 06/07/76		UK N	٠.					•	•	•	232 233
2047 2047	•	06/04/76 06/10/76				9 4.51 1 4.69				:	:	:	2 14
2047	•	10/03/17		UK	•	•	•	•	•	•	•	•	235 236
2051 2242	•	06/07/76 06/07/76			•	:	:	:	:	:	:	:	237
2249		06/08/76	. N T	UK		•	•	•	•	•	•	•	238 239
2252 2252		09 21 .6 A 06/07/76 06/11/76				2 6.08	• :	4.2	:	:	:	:	240
2252	•	06/12/76	. !	51 N	в.	6.19	9.	•	•	•	•	•	241 242
2252 2252		09 22 45 0 06/07/76 06/09/76				8 2.64				:	:	:	243
2252	•	06/09/76		UK 5	2.		2.0	7.	•	•	•	•	244 245
2252 2253		06/10/76 06/09/76			8 4.0	6 3.09	5 2.1	2 1.8	•	:	:	:	246
2256		05/29/71		UK									247

GLO	POSITION (1950)	DATE (UT)	SAO# COMM	1,65	2.28	3.2	3.5	4,9	8.4	11,2	12.5	CRD
2265	•	05/29/78	NP UK	•				•				248
2271 2277	:	05/29/78 06/07/76	NF UK	:	:	:	:	:	:	:	:	249 250
2290	18 56 04. +06 38 .3	A 05/29/18	P OK		4.15		0.72	-0.72	-2.51		-3.44	251
2298 2303	•	05/30/78 05/30/78	NP UK NP UK	:	•	:	:	•	•	•	•	252 253
2438 2441	•	11/19/76	NP UK	•	•	•	•	·	:	:	:	254
2442	•	06/0//76	NP UK NP UK	•	:	:	:	:	:	:	:	255 256
2445 2445	19 42 21. +35 07 .9	P 06/07/76 10/03/77	ያ UK ያ UK	•	•	•	•	•	•	•	•	257
2448 2455	•	06/07/76	NP UK		•	:	:	:	:	:	:	258 259
	•	06/08/76	NP OR	•	•	•	•	•	•	•	•	260
2455 2457	•	10/16/76 05/30/78	NP UK NP UK	:	:	•	•	•	•	•	•	261 262
2470 2474	10 53 86 .33 18 1	06/07/76	NF UK	•	••••	•				:		263
2477	19 53 46. +22 14 .1 19 54 50.0 +30 35 57	1 05/30/78 U 12/10/76	P OK P UK	:	2.93 8.66	:	1.99	1.14	0.14	-1,97	-0.40 -2.23	264 265
2478 2481	•	10/16/76 06/07/76	ክም ሆኖ ያ ጽ	•	•	•	•	•	•	•	•	266 267
2481	19 55 55.0 -03 41 24	5 06/11/76	143959 K	2.25	2.03	1.89	1.84	:	:	:	:	268
2489 2492	•	10/16/76 10/16/76	NP UK	:	:	•	:	:	:	:	•	269 270
2494	19 59 21. +40 45 .7	P 06/07/76	P OK	_								271
2494	•	10/03/77	P UK					:	:	:	:	272
2636 2536	20 40 47.0 +42 45 52	U 06/11/76 06/12/76	P UK NY P NY NB	8.76	6.97 7.07	5.61	5.34	:	•	:	:	273 274
2636 2636	•	11/20/76	NF UK PUK	•	•	•	•	•	•	•		275
2636	:	12/11/76	P UK	• • • • • • • • • • • • • • • • • • • •	•	:	6.69	5.12	2.99	2.56	2.63	276 277
2636 2638	•	10/03/77 06/07/76	PUK NPUK EO	8.95	7.18	:	5.53	:	:	:	•	278 279
2669		12/10/76	NP UK	•	•	•	•	•	•	•	•	240
2679 2679	20 54 55.8 +37 13 35	10/11//6	P UK NB	5.17	3.73	:	2.25	•	•	:		251 262
2679 2681	•	11/18/76 06/08/76	F UK NP UK	•	3.67	•	2.12	1.34	0.40	-0.12	•	283 263
2686	20 57 00.7 +27 14 42	U 10/11/76	P UK	5.51	3.00	:	0.51	:	:	•	:	245
2686 2789	21 38 12. +50 00 .8	11/18/76 ¥ 12/16/76	P UK P K N8	:	2.93	:	0.36	0.94	-2.24	3.0H -0.9H	- 1, 4£	28≀ 2∺1
2789	•	10/24/77	P K	7.15	4.81	•	•		•	•	•	288
2824 2932	22 38 34. +49 45 .6	06/08/76 ¥ 10/12/76	NP UK P K	2.26	•	:	:	:	:	:	:	2 H P 2 9 0
2932	22 38 34. +49 45 .6	V 11/18/76	P K	•	1.68		1.32	1.38	0.93	-0.04		291
2939 2940	22 40 37.0 +27 53 42	10/11/76 \$ 12/05/76	NP UK 090732 K	:	1. 19	•	0.89	1.02	0.77	0.24	0.36	292 293
2941		I 12/10/76	₽ K	•	2.24	•	1.48	1, 19	0.14	-0.64	-0.51	294
2944 2954	:	10/16/76 12/10/76	NP UK	:	:	:	:	:	•	:	:	295 296
2961 2999	•	12/10/76 10/15/76	NP UK NP UK	:	:	:	:	•	:	•	•	297 298
3008	•	10/15/76	NP UK	•	•						·	299
3139	•	10/15/76	NF OK	•	•	•	•	•	•	•	•	300
3140 3143		S 12/05/76 V 12/10/76	053355 K	:	1.88	•	1.55	1.78	1.62	1,26	1,16	301 302
3144 3144	•	10/11/76 10/16/76	NP UK NP UK	•	•	•	•	•	•	•	•	30 I 30 4
3151	•	10/16/76	NF UK	:	:	:	:	:	:	:	:	30%
3178 3196	23 58 41.9 +60 04 37	12/11/76 5 12/10/76	NF UK 021002 K	:	0.54	:	0.23	0.36	0.00	-0.16	-0.03	306 307
4009 4010	•	10/11/76	MP UK MP UK	•	•	•	•	•	•	•	•	30 H 30 P
4012	•	10/11/76 10/11/76	NP UK	:	:	:	:	:	•		:	310
4013	01 52 47.6 +16 56 41	s 10/16/76	092697 K	2.20	1.90		1.93					111
4013 4029	02 57 32.5 +60 17 22	11/18/76	092697 K	9.36	1.91	•	1.61	1.70	0.90	0.00		112
4029	*	10/15/76	P K NB		7.22	:	•	:	·	÷	:	314
4029 4029	•	11/18/76	NP K NP K	:	:	:	•	•	:	:	:	315 316
4024	•	12/16/76	7 K NB	•	6.96	•	4.76	3.47	1,94	1.61	1.09	317
4029 4029	•	10/03/17	P K	10.06	9.01	:	:	:	:	:	:	31H 11+
4029 4029	•	10/25/77 10/25/77	PK PK	9.37	6.92	•	•	•	•	•	٠	120 341
4029	•	10/24/77	7 K	:		:		0.76		1.41	•	3.2
4029 4036	•	10/26/17	F R NF UK	:	8.84	•	7.07	4.5H	2.34	2.12	1.76	123 320
4038	•	10/11/76	NF UK N♥ UX	•	•	•	•	•	•	•	•	325
4040	•	10/11/76	S Z UK	:	•	:	•	:		:	:	926 327
96 69 96 69	•	02/08/11	HE DE ED	•	:		:		•		:	129
4139	11 52 01. +17 25 .2	I 05/11/7H	PR	÷	2.95	:	7.62	2	:	:		; ;

GL•	POSITION	(1950) EATE (UT	SAOF CONN	1.65	2.28	3.2	3.5	4.9	8.4	11.2	12.5	CRD
4139	11 62 19 1	1 +37 02 37 5 02/U8/77	062754 K		2.17		1.93	1.98	1.69	1.72	1.87	331
4166	11 12 376 3	02/22/78					•		•	•	•	112
4184	•	02/21/78						•	•	•	•	333
4184	•	05/27/76			•		•	. •		•	•••	334
4219	15 46 30.7	7 +28 18 32 S 05/27/78	084015 K		4.21	•	2.95	1.63	0.06	•	-0.10	335
4253		7 +09 20 39 U 10/16/76		9.42	7.04	•	4.17	•		•••	•	336 337
4253		11/20/76	F UK	•	6.64	•	•	2.54	•	0.1A	•	337
4254		10/16/76		•	•	•	•	•	•	•	•	339
4255		05/30/78		•	•	•	•	•	•	•	•	340
4257	•	10/16/76	NP UK	•	•	•	•	•	•	•	•	,40
		11/20/76	NF UK			_	_					341
4257 4295	22 59 37.			:	3.04	:	1.98	1.18	-0.04	-0.95	-0.53	342

Position codes (sources of quoted positions)

Λ	AFGL (no finding chart available) ($\alpha \pm \beta'$, $\delta \pm 2$).
В	Catalogue of Bright Stars (Schlesinger and
	Jenkins 1940) ($\alpha \pm 1'$, $\delta \pm 1'$).
1	AFG1 (see finding chart, Fig. 5) ($\alpha \pm 3'$, $\delta \pm 2'$)
l .	Two-micron sky survey (IRC) (Neugebauer and Leighton (1969) ($\alpha \pm 1'$, $\delta \pm 3'$)
S	SAO star catalog (SAO 1966) ($\alpha \pm 0.5$ ", $\delta \pm 0.5$ ").
U	Univ. California, San Diego (positions measured from Palomar Sky Survey plates, see also finding charts in Figs. 4 and 5) ($\alpha \pm 10^{\circ}$, $\delta \pm$ 30°)
V	General catalog of variable stars (Kukarkin et al. 1969) ($\alpha \pm 1'$, $\delta \pm 1'$)

Comment codes K Known or identified source (ID given by AFGL) UK Unknown or unidentified source (no ID given by AFGL)

AIGL)

EO Extended object (with respect to AFGL beam size).

NV Not visible (in Mt. Lemmon 1.5 m).

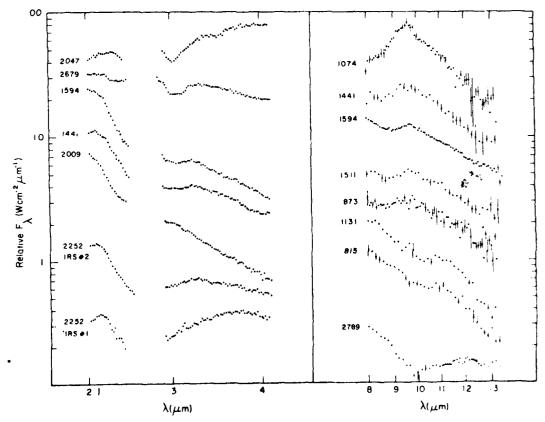
S1 Source No. 1, etc. (if more than one in error box).

Not found in standard error box

Found in error box

Found in error box
NB Narrow-band observations taken
A Following magnitudes, indicates an approximate value

Note GL 2474 does not appear in the GL Catalog, and therefore is more appropriately identified at CRL 2474.



Eur. 2 - 2 - 4) and 8 - 14-μm spectrophotometry plotted by channel, with indicated wavelengths intended only as a rough guide. The data have been multiplied by the following factors for each object - 2047, 8 4716, 1074, 3 6516, 2679, 2,4616, 1441, 1 1617, 1594, 7 6133, 1594, 8 2815, 1441, 2 3615, 1511, 2 8816, 2009, 1 2816, 873, 2 2415, 22523, 4 6914, 1131, 4 3015, 22521, 2 1415, 815, 4 4715

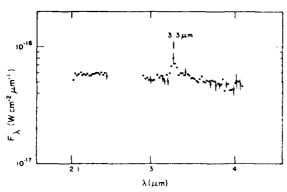


Fig. 3. The 2-4- μ m narrow-band spectrum of GL 2636 plotted by channel, with indicated wavelengths intended only as a rough guide. The feature at 3.3 μ m, originally discovered in the planetary nebula NGC 7027 (Merrill, Soifer, and Russell 1975), occurs in a limited class of infrared sources.

limits, where the latter were taken at the AFGL's claimed levels of statistical completeness. Although some fainter objects are listed in the catalog, the AFGL detection limit at 11 μ m suggests that for an object detected by the AFGL at 11 μ m, and with a blackbody spectrum warmer than 320 K, a detection at 3.5 μ m with the InSb system would be more likely than a detection at 11.2 μ m. Correspondingly, for a scan at 3.5 μ m with the Si(As) system, the limiting blackbody temperature would be 410 K.

III DISCUSSION

Table II presents complete results to date for the broad-band photometry. Some of these data also appear in Fig. 1, with the remaining data either insufficient for classification or showing featureless zero-color (K - N= 0.0) stars. Figure 1(a) shows those objects having 2.3 12.5- μ m color temperatures < 1000 K. Some of these spectra show a small depression at 5 μ m likely to be due to CO, or emission or absorption at 11.2 µm likely to be due to SiC or silicates. Figures 1(b) and 1(c) show those objects classified as relatively warm ($\Gamma > 1000 \text{ K}$) and as having silicate or SiC emission. The strengths of the features in Fig. 1(b) were judged to be stronger than those of Fig. 1(c). Finally, Fig. 1(d) shows objects with unusual characteristics, subject to confirming measurements. In particular, GL 371 apparently shows a remarkable absorption feature in the region of 5 µm.

From the broad-band photometry, we infer that sources 276, 416, 538, and 2636 may be multiple sources. For 276 and 416, this results from discrepancies in the photometry in Table 1. For 538 and 2636, Fig. 1 shows striking discrepancies between AFGL data and UCSD data, which also may be due to the large aperture (3.4) used by the AFGL and the possible extended nature of these sources. K. M. Merrill (1978, private communication), however, has confirmed the multiple nature of 2636. GL 2023, 2252, and 4029 were found to be multiple sources by direct recordings during the raster scans, with individual sources to be hereafter designated as Nos.

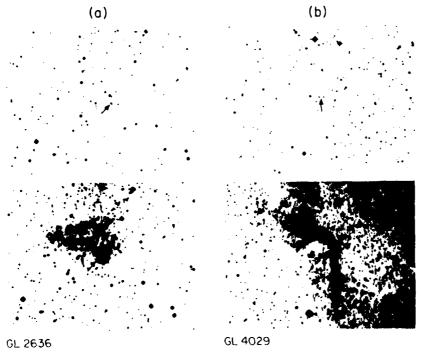
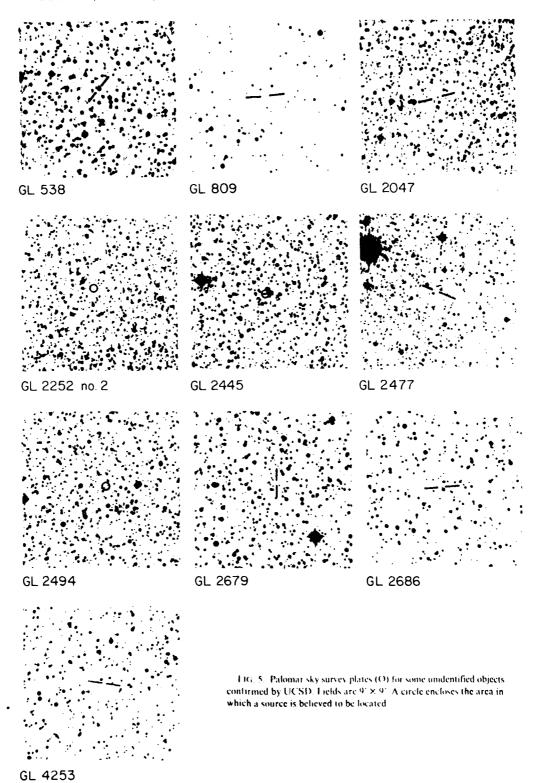


FIG. 4. (a) O (top) and F Palomar sky survey plates (north is up) for GL 2636. B supergiant is indicated, and the near-infrared source is situated 15". W and 15". N of that location 1950 coordinates for the infrared source are $\alpha = 20940^{\rm m}47^{\rm e}0 (\pm 10^{\rm e})$, $\delta = \pm 42^{\rm e}45^{\rm e}52^{\rm e} (\pm 10^{\rm e})$. Fields are $17' \times 13''$ (b) O (top) and F Palomar sky survey plates (north is up) for GL 4029. An infrared source is indicated between what are apparently reflection nebulae. 1950 coordinates are $\alpha = 102^{\rm h} (57^{\rm m} - 32^{\rm m} (\pm 10^{\rm e}))$. $\delta = \pm (610^{\rm e}) 7 (22'' (\pm 10'')$. Fields are $17' \times 13''$.



1, 2, etc., in the order of detection during scanning of the error box.

547

The narrow-band spectrophotometry appears in Figs. 2 and 3. From these data we have made the following tentative classifications: GL 2679 (this work and Cohen and Kuhi 1976) and 2047 are carbon rich; GL 815 and 1131 are probably carbon rich; GL 1074, 1441, 1594, 2009, 2252 No. 1, and 2252 No. 2 are oxygen rich; and GL 873 and 1511 are probably oxygen rich.

Noteworthy objects for which narrow-band spectra have been obtained are GL 2636 and GL 4029. The spectrum of GL 2636 (Fig. 3) possesses the unidentified 3.3- μ m feature discovered by Merrill, Soifer, and Russell (1975) in the planetary nebula NGC 7027. GL 2636 appears to be a multiple source, with both a near- and a far-infrared source in close proximity (30") to a B supergiant (Merrill 1978, private communication). The spectrum in Fig. 3 is of the near-infrared source, which is located 15" W and 15" N of the B supergiant indicated on the finding chart of Fig. 4(a). 1950 coordinates for this object are $\alpha = 20^h 40^m 47!0 (\pm 10")$, $\delta = \pm 42^o 45' 52" (\pm 10")$.

GL 4029 is also a multiple source, where two compact sources have been found within a region of extended infrared emission. The following spectral features have been found in at least one of the sources: in absorption, the 3.1-µm ice feature, and in emission, the unidentified

3.3-, 6.2-, 7.7-, and 11.2- μ m features. In addition, based on optical spectra obtained at Lick Observatory, the two compact sources are situated on or near two optical objects that appear to be reflection nebulae. The infrared spectra of the compact sources are not presented here due to some remaining confusion in determining to which of the several objects in this region the data pertain. A point between the optical objects is indicated in Fig. 4(b). The 1950 coordinates for one of the infrared sources are $\alpha = 02^{\rm h}/57^{\rm m}/32^{\rm s}/5 ~(\pm 10^{\rm s})$, $\delta = +60^{\rm o}/17^{\rm s}/22^{\rm s}/(\pm 10^{\rm s})$. Finding charts for additional objects appear in Fig. 5.

At present, results concerning the statistical nature of the objects in the AFGL catalog would be biased if based on the data presented here. Furthermore, any interpretation of the fraction of multiple sources found so far would be difficult to make since the sources were detected at flux levels below the limits of the AIGL Future observations, however, are being designed to admit valid statistical conclusions about the AFGL sources.

Support for these observations was provided by AFGI contract No. F19628-76-C-0252. We would like to thank C. A. Lindsey, R. W. Russell, B. T. Soifer, and S. P. Willner for advice and for assistance with instrumentation and observations.

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1. Table II contains three minor errors:

AFGL 0873: add the comment NB; 3.5 µ mag = 1.73 instead of 11.73.

AFGL 2686: 4.9 µ mag = -0.84 instead of 0.84; 11.2 µ mag = -3.08 instead of 3.0:

AFGL 3140: identified with SAO 053335, instead of 053355; coordinates 23^h 42^m 10.6^s, +41° 46′ 52″.

- 2. The GL designation is slightly confusing, since the Gliese catalog already uses this indicator (R.S. Harrington, personal communication). We suggest referring to the infrared survey objects by the full prefix AFGL until IAU nomenclature is established.
- 3. Table II lumps together objects reported in the early AFCRL survey (Walker and Price, 1975) and in the later, improved AFGL survey (Price and Walker, 1976). The following list of 102 objects in Table 1 appear only in the AFCRL catalog. These have a higher rate of non-confirmation of unidentified sources than do the objects in the current AFGL catalog. Of the sources only in the AFCRL catalog, only three (AFCRL 538, AFCRL 1671, and AFCRL 2474) were "found". These are likely to be accidental discoveries since their magnitudes are fainter than the AFGL detection limits. We list them here only to be able to present our complete data.

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Walker, S.P., and Price, R.P. (1975), "AFCRL Infrared Sky Survey,"

AFCRL TR-75-03.

Amendment to Table II
OBJECTS APPEARING IN THE AFCRL CATALOG

105	1180	1507	1645	1729	2022
119	1182	1518	1646	1730	2249
145	1193	1520	1647	1731	2253
538	1194	1522	1649	1735	2265
1005	1197	1524	1655	1737	2277
1006	1290	1528	1657	1738	2298
1011	1297	1529	1662	1739	2438
1013	1305	1530	1664	1742	2441
1015	1373	1573	1665	1795	2457
1019	1375	1574	1667	1802	2470
1029	1377	1577	1668	1808	2474
1031	1382	1578	1671	1812	2478
1032	1383	1580	1672	1813	2 4 8 9
1046	1384	1582	1674	1815	2638
1048	1385	1587	1723	1884	2669
1049	1390	1590	1725	1897	2944
1054	1391	1644	1727	1901	2954